

A peculiar Of star in the Local Group galaxy IC 1613[★]

A. Herrero^{1,2}, M. Garcia^{1,2}, J. Puls³, K. Uytterhoeven^{1,2}, F. Najarro⁴, D.J. Lennon⁵ and J.G. Rivero-González³

¹ Instituto de Astrofísica de Canarias, C/ Vía Láctea s/n, E-38200 La Laguna, Tenerife, Spain.

² Departamento de Astrofísica, Universidad de La Laguna, Avda. Astrofísico Francisco Sánchez s/n, E-38071 La Laguna, Tenerife, Spain.

³ Universitätssternwarte München, Scheinerstr. 1, 81679 München, Germany.

⁴ Centro de Astrobiología (CSIC-INTA), Ctra. de Torrejón a Ajalvir km-4, E-28850, Torrejón de Ardoz, Madrid, Spain.

⁵ ESA, Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA.

ABSTRACT

Context. Results from the theory of radiatively driven winds are nowadays incorporated in stellar evolutionary and population synthesis models, and are used in our interpretation of the observations of the deep Universe. Yet, the theory has been confirmed only until Small Magellanic Cloud (SMC) metallicities. Observations and analyses of O-stars at lower metallicities are difficult, but much needed to prove the theory.

Aims. We have observed GHV-62024, an O6.5 IIIf star in the low-metallicity galaxy IC 1613 ($Z \approx 0.15Z_{\odot}$) to study its evolution and wind. According to a previous preliminary analysis that was subject to significant restrictions this star could challenge the radiatively driven wind theory at low metallicities. Here we present a complete analysis of this star.

Methods. Our observations were obtained with VIMOS at VLT, at $R \approx 2000$ and covered approximately between 4000 and 7000 Å. The observations were analysed using the latest version of the model atmosphere code FASTWIND, which includes the possibility of calculating the N III spectrum

Results. We obtain the stellar parameters and conclude that the star follows the average wind momentum–luminosity relationship (WLR) expected for its metallicity, but with a high value for the exponent of the wind velocity law, β . Comparing this with values of other stars in the literature, we suggest that this high value may be reached because GHV-62024 could be a fast rotator seen at a low inclination angle. We also suggest that this could favour the appearance of the spectral "F"-characteristics. While the derived β value does not change by adopting a lower wind terminal velocity, we show that a wrong V_{∞} has a clear impact on the position of the star in the WLR diagram. The N and He abundances are very high, consistent with strong CNO mixing that could have been caused by the fast rotation, although we cannot discard a different origin with present data. Stellar evolutionary model predictions are consistent with the star being still a fast rotator. We find again the well-known mass-discrepancy for this star.

Conclusions. We conclude that the star follows the WLR expected for its metallicity. The results are consistent with GHV-62024 being a fast rotator seen close to pole-on, strongly contaminated at the surface with CNO products and with a wind structure altered by the fast rotation but without modifying the global WLR. We suggest that this could be a general property of fast rotators.

Key words. Galaxies: individual: IC1613 – Stars: early-type – Stars: fundamental parameters – Stars: mass-loss – Stars: rotation – Stars: evolution

1. Introduction

Three main parameters determine the structure and evolution of a star: its mass, angular momentum and metallicity. For massive stars the first two parameters change rapidly during their lives because of stellar winds, whose strength is determined by the third parameter. A reliable theory of stellar winds for massive stars is accordingly crucial for our predictions about how these stars evolve and therefore for our predictions of their properties (ionizing flux, mechanical energy delivered, chemical composition, etc.) at a given moment in their lives and final fate (type of supernova and remnant left behind).

The theory of radiatively driven winds, used to explain the observed winds in massive stars and whose predictions are incorporated in stellar evolutionary and population synthesis models, is based on the work by Castor, Abbott & Klein (1975), with improvements by Pauldrach, Puls & Kudritzki (1986) and Friend & Abbott (1986). Recent reviews can be found in

Kudritzki & Puls (2000) and Puls, Vink & Najarro (2008). One of the most important predictions of the theory is that the modified wind momentum rate of the stellar wind (MWM, the product of mass-loss rate, wind terminal velocity and square root of the stellar radius, which we denote as $D_{\text{mom}} = \dot{M}v_{\infty}R^{1/2}$ following the use in the literature) should depend almost solely on the stellar luminosity and metallicity. The VLT-FLAMES survey of massive stars (Evans et al. 2005, P.I. S. Smartt) was a project aimed precisely at testing this prediction, which was nicely confirmed by Mokiev et al. (2007b) by analysing O-type stars in the Milky Way, the LMC and the SMC. However, the second part of the project, the comparison of the predictions of stellar evolutionary models including mass-loss and rotation with the stellar properties, revealed problems when comparing the nitrogen abundances (Hunter et al. 2008), indicating that other physical processes may be at work. This started the VLT Tarantula FLAMES Survey (VFTS), a detailed investigation of the massive stellar population of 30 Doradus in the LMC currently under way (Evans et al. 2011). There a deeper study of the behaviour of the LMC O-stars winds will be possible, but without extending the available metallicity baseline.

Send offprint requests to: A. Herrero (e-mail: ahd@iac.es)

[★] Based on observations obtained at the ESO VLT for Programmes 078.D-0767 and 080.D-0423.

Very many O-stars have been analysed in the Magellanic Clouds (see, among others, Bouret et al. 2003; Hillier et al. 2003; Mokiem et al. 2006; Mokiem et al. 2007; Massey et al. 2004 or Massey et al. 2005). However, going to lower metallicities implies a large distance jump. Yet, many star-forming galaxies in the Local Universe and presumably most stars in the early epochs of the Universe have metallicities below that of the Small Magellanic Cloud. It is therefore natural to try to go one step further and analyse stars at even lower metallicities, studying their wind behaviour and stellar evolution. Although going to galaxies farther away and with lower metallicity than the SMC will result in an increasing difficulty to obtain good quality spectra that can be analysed, present day telescopes and instruments clearly allow for the next step. At metallicities of about half that of the SMC and distances of one to a few megaparsecs it is still possible to perform reliable analyses. This has been done using the visually brightest objects, the A- and B- supergiants, by Urbaneja et al. (2008) in the WLM (0.96 Mpc), Evans et al. (2007) in NGC 3109 (1.3 Mpc), Castro et al. (2012) in NGC 55 (2.0 Mpc) or Urbaneja et al. (2005) in NGC 300 (2.2 Mpc). Bresolin et al. (2001) have even analysed individual A supergiants in NGC 3621 at 6.7 Mpc. The hotter O-stars, however, are visually fainter and much more difficult to observe at far distances. For this reason, there have been no spectroscopic analyses of individual O-stars located in low-Z galaxies using modern standard techniques (i.e., analysing optical spectra by means of NLTE, spherical models with mass-loss), not even in the Local Group, until very recently (Herrero et al. 2011; Trammer et al. 2011).

IC 1613 is a dwarf irregular galaxy in the Local Group with a distance modulus of $(m - M)_0 = 24.31 \pm 0.06$ (Dolphin et al. 2001). It has a metallicity of $\log(O/H) + 12 = 7.80 \pm 0.10$ as determined from its B-supergiants (Bresolin et al. 2007), whereas nebular studies vary between 7.60 and 7.90 (Lee et al. 2003). The metallicity of the B-supergiants indicates $Z_{\text{IC 1613}} = 0.13 \pm 0.07 Z_{\odot}$ (using solar abundances from Asplund et al. 2009). This value is below the value derived by most authors for the SMC O- and B-stars. For example, Heap et al. (2006) obtain $Z_{\text{SMC}} = 0.2 Z_{\odot}$ from an analysis of O-stars and Dufton et al. (2005) obtain $Z_{\text{SMC}} = 0.25 Z_{\odot}$ from an analysis of B-supergiants. The situation is similar for other authors, e.g. Bouret et al. (2003), who obtained $Z_{\text{SMC}} = 0.2 Z_{\odot}$. Therefore the stellar metallicity in IC 1613 is below the SMC value and the study of its stars represents a step forward in our analyses at low metallicities.

IC 1613 shows a recent and intense burst of massive star formation, particularly in its NE part. We have recently published a new catalogue of OB associations in IC 1613 (Garcia et al. 2009, identified as [GHV2009]) and their physical properties (Garcia et al. 2010) as part of our effort to carry out an in-depth study of the young population of IC 1613.

We also obtained spectra of some stars in this galaxy using VIMOS at VLT. In the field of IC 1613 we found the star 62024 in the GHV2009 catalogue (hereafter we identify the star as GHV-62024), an Of star. This star is ideally suited to study the stellar evolution and winds in IC 1613, because it displays strong N III lines and a P-Cygni profile in He II 4686. A preliminary analysis of this star by Herrero et al. (2011) resulted in a too high $\log D_{\text{mom}}$ as expected from its metallicity and luminosity, which would challenge the theory of radiatively driven winds at low metallicities. An even more serious challenge is the very recent work by Trammer et al. (2011). These authors analysed six O-stars in low-Z galaxies (four in IC 1613 and one in WLM and NGC 3109). The four stars in IC 1613 tend to show too high values of D_{mom} , although they are not inconsistent with

the theory because of the large error bars. But the two other stars have too high D_{mom} values, not consistent with theoretical predictions. However, although these two recent works seem to point towards a problem between observations and theory, we note that the analyses are subject to strong assumptions.

In the present work, we carry out a detailed analysis of GHV-62024, showing that when we drop the assumptions made by Herrero et al. (2011), we obtain a MWM value consistent with the theoretical expectations, although the analysis opens alternative questions. The observations and data reduction are presented in Sect. 2. The spectrum is described in Sect. 3. The spectral analysis follows in Sect. 4 and a discussion of the stellar wind momentum rate is given in Sect. 5. The evolutionary status and other properties are discussed in Sect. 6. Finally, the conclusions are presented in Sect. 7.

2. Observations and data reduction

Details of the observations are given in Herrero et al. (2010). We repeat here the main points. Observations were performed with VIMOS at VLT in MOS (multi-object spectroscopy) mode. The HR-blue and HR-orange gratings were used, resulting in a resolution of $R \sim 2050$ and $R \sim 2150$, respectively, and a spectral wavelength coverage roughly from 3870 to 7240 Å (depending on the star position). We had 19 observing blocks for the blue part of the spectrum, each with three individual exposures, and ten observing blocks for the red part, each with two individual exposures. Thus, in total 3×19 blue spectra were observed from October 5 to November 5, 2007, and 2×10 red spectra from November 5 to November 13, 2007. Integration times for individual exposures were 650 sec for the blue spectra and 1005 to 1110 sec for the red spectra. The three (two) consecutive exposures in the blue (red) were coadded. The coadded spectra were extracted with standard IRAF¹ procedures after wavelength calibration and cosmic ray removal. Next, a barycentric correction and a correction for the systemic velocity of IC 1613 ($-234 \pm 1 \text{ km s}^{-1}$, Lu et al. 1993) were applied. Finally, after checking that no significant radial velocity variations are present, all 19 (10) resulting spectra were coadded to a single blue (red) spectrum. The resulting spectra, with an S/N of 85 both in the blue and red wavelength regions, were then rectified.

Figure 1 shows the finding chart of GHV-62024. The star is located at $\alpha(2000) = 01^{\text{h}} 05^{\text{m}} 00.72^{\text{s}}$, $\delta(2000) = 2^{\circ} 8' 49''.1$ (from the astrometry by Garcia et al. 2009). It clearly lies in a zone of intense recent star formation, with strong ionized bubbles visible in Fig. 1. Table 1 gives the optical photometric data of the star from the catalogue of Garcia et al. (2009).

Table 1. Photometric data for GHV-62024 from Garcia et al. (2009).

U	B	V	R	I
18.69	19.44	19.60	19.76	19.68

¹ IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy, Inc., under cooperative agreement with the National Science Foundation.

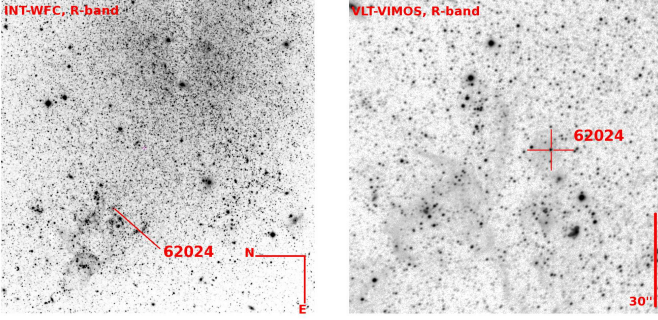


Fig. 1. Finding chart for GHV-62024. The left part is an image obtained with the WFC at the INT in the R band. The region imaged (5x5 arcmin) contains the IC 1613 intense star-forming region, with many ionized bubbles. The right-hand side is a VLT-VIMOS image in the same band, at a larger scale (indicated by the scale bar at the right edge). The orientation is the same in both images.

3. Description of the spectrum

Figures 2 and 3 show the observed coadded blue and red spectra of GHV-62024, respectively (together with the best-fit model that we describe below). The stellar spectrum is dominated by the H and He lines, plus a conspicuous N III emission at 4630–40 (but without N IV) and a weak Si IV 4089 (but Si IV 4116 could not be detected). The core of the Balmer lines show an intense nebular contamination, but fortunately the wings are well defined (except for H α). The red spectrum has a poorer quality, and no individual features could be unambiguously identified, apart from H and He lines.

The star has a He II 4541 line stronger than He I 4471. By measuring the equivalent widths of these lines ($\log(W') = \log(EW_{4471}) - \log(EW_{4541}) = -0.18$, see Conti & Alschuler 1971; Markova et al. 2011), and comparing with standards we classify the star as O6.5. Because of its N III 4634–40 and He II 4686 emissions we assign it the f-class. The luminosity criteria of O stars are based on metal lines that are very faint in our spectrum and are usually defined at other metallicities. We assign the star a luminosity class based on the strength of the Balmer lines. Thus the star is classified as O6.5 IIIf. The red spectrum is heavily contaminated by nebular and sky emission and consequently H α yields no information on mass-loss rate. Therefore, the P-Cygni profile at He II 4686 is a very important signature. The 19 individual spectra have an average SNR of 20 (with a large scatter). We then inspected the spectra of individual observing blocks to ensure that the P-Cygni profile is not caused by unidentified cosmic rays, spikes, or just random noise variations. We also inspected the nebular emission around the star, and confirmed that there is no emission in He II 4686. We conclude that the P-Cygni profile is a stellar feature.

4. Spectral analysis

We analysed GHV-62024 with the newest FASTWIND version, which includes dielectronic recombination (Rivero-González et al. 2011a). The analysis was based mainly on H δ , H γ and H β , He I 4471, He II 4541, 4686 and the N III emission lines at 4634, 4640 Å. Although all line profiles are sensitive to all parameters, the main diagnostics for gravity was the wings of the low Balmer lines (except H α which was not used). For the temperature we mainly used the ratio He I

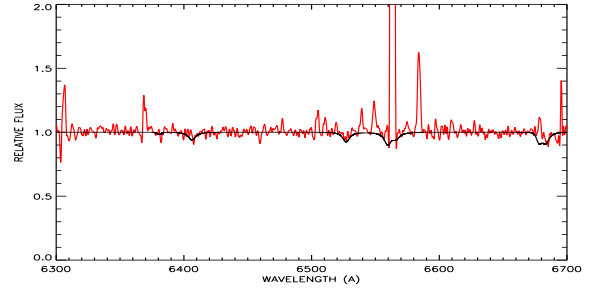


Fig. 3. Same as Fig. 2, red spectrum.

4471 to He II 4541, giving some weight also to the He I 4387, 4922 lines (which point to a temperature slightly cooler than He I 4471). The ratio of N III 4634–40 to N IV 4058 also was a secondary indicator (the N IV line is not present, thus giving an upper limit for the temperature). Without a useful H α profile, the mass-loss rate and the exponent of the velocity law, β , were determined using the He II 4686 line. This is crucial, because there is a degeneracy between the β exponent and the derived mass-loss rate. In Fig. 4 we see the fit to the He II 4686 profile using the value of $\beta = 2.0$ and $\beta = 1.0$. From the figure it is clear that the higher value fits the observed profile better. It turned out to be impossible to find a good fit to the red wing of He II 4686 with a low β value. The abundances of N and He, given in Table 2, were fixed by fitting the profiles of all lines. The first one is given in the usual logarithmic scale in which the H abundance is 12, while the He abundance is given by number, by $Y_{\text{He}} = (N_{\text{He}}/N_{\text{H}})$. The red spectrum was not used for the analysis, but we note that it is consistent with the adopted parameters (see Fig. 3).

The final stellar parameters are given in Table 2 and the best fits are shown in Figs. 2 and 3. For the analysis we adopted a projected rotational velocity of 80 km s $^{-1}$ and a radial-tangential macroturbulent velocity of 50 km s $^{-1}$. Although we were unable to determine these values from the line profiles, at a spectral resolution of $R \approx 2000$ they have little impact on the parameter determination, because the line broadening is limited by the spectral resolution or the Stark broadening. The adopted metallicity for the model calculations was $Z = 0.15 Z_{\odot}$ at the beginning and we changed to $Z = 0.13 Z_{\odot}$ for the models close to the final stellar parameters (adopting the metallicity obtained by Bresolin et al. 2007 for B-supergiants, although this difference in Z produced no appreciable change in the spectrum). Unfortunately, there are no features in our spectrum that allow us an independent metallicity determination. The radius was determined by fitting the visual magnitude by integrating the flux of our final model in the Johnson V-filter and after correcting for the distance to IC 1613 and the reddening (see Herrero et al. 1992 or Repolust et al. 2004). We used the photometry from Garcia et al. (2009), quoted in Tab. 1. The final model has an intrinsic $(B - V)_0 = -0.26$, while the observed one is $(B - V) = -0.16$. Assuming a standard extinction law with $R_V = 3.1$ and with the adopted distance modulus of 24.31 we finally obtained an absolute magnitude $M_V = -5.03$, which results in a radius $R = 11.1 R_{\odot}$ and using the derived effective temperature, in a luminosity $\log(L/L_{\odot}) = 5.29$. These values are consistent with a slightly underluminous O6.5 giant, which reduces the possibility that the star is actually a binary system with two comparable components that could affect the spectral analysis.

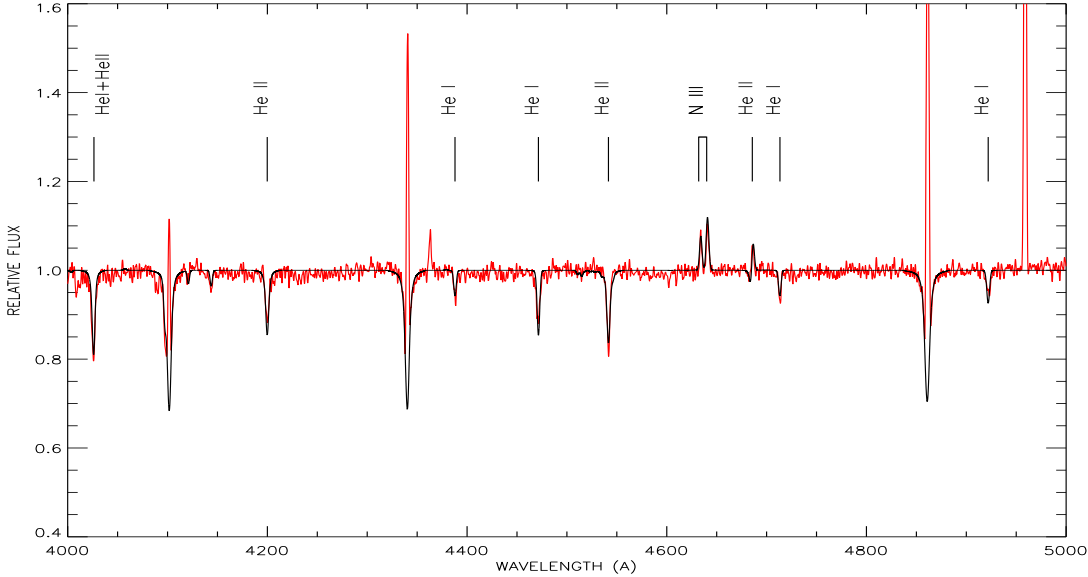


Fig. 2. Observed blue spectrum of GHV-62024 (red) and the adopted best fit (black). The main He and N lines used in the analysis have been marked. Stellar parameters are given in Table 2. The emission in the red wing of H_γ is a sky line.

We have little information on the terminal velocity. We initially adopted $V_\infty/V_{\text{esc}} = 3.4$ and $V_\infty(Z) \propto Z^{0.12}$ as given by Castro et al. (2012) (see also Castro 2010), who used data from the literature (the metallicity dependence has been taken from Leitherer et al. 1992). We note that this ratio is higher than the more commonly used one of 2.6, which would result in a terminal velocity 30% lower. This has little effect on the analysis, but may be important for the MWM (see next section). We adopted a wind terminal velocity of 1800 km s^{-1} . This is consistent with the absorption part of the He II 4686 P-Cygni profile, which indicates a lower limit of $\approx 1200 \text{ km s}^{-1}$, although a more precise determination of V_∞ from this line is not possible. The mass-loss rate is derived mainly from the emission profile of He II 4686. We obtain $\dot{M} = 4.3 \times 10^{-7} M_\odot \text{ yr}^{-1}$.

We estimated the error in T_{eff} to be $\pm 1500 \text{ K}$ and in $\log g \pm 0.15$ dex, which translates into an uncertainty of ± 0.6 in R/R_\odot and ± 0.09 in $\log(L/L_\odot)$. The spectroscopic stellar mass (the mass obtained using the values of $\log g$ and radius obtained from the analysis of the spectrum, and corrected from the centrifugal force) results in $M_{\text{sp}} = 18.3_{-5.5}^{+7.8} M_\odot$. The error in the helium abundance is estimated to be ± 0.04 and in $\log(N/H) \pm 0.30$ dex. The uncertainty in V_∞ is large, $\pm 600 \text{ km s}^{-1}$. The uncertainties in β and \dot{M} are more difficult to estimate because the upper limit of β (and thus the lower limit of \dot{M}) is not well constrained. However, if we consider β fixed, the uncertainty in \dot{M} (which is then a lower limit) is ± 0.30 dex.

Table 3. Synthetic photometry for the model given in Table 2. Values are absolute magnitudes.

U	B	V	R	I
-6.37	-5.29	-5.03	-4.84	-4.61

5. The wind momentum of GHV-62024

The first remarkable result of the previous section is the high value we obtain for β , the exponent of the wind velocity law.

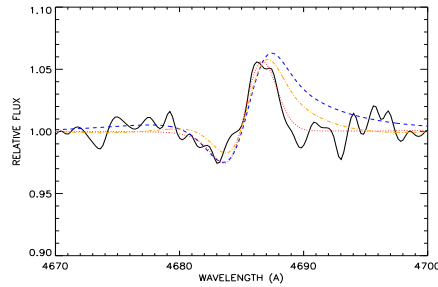
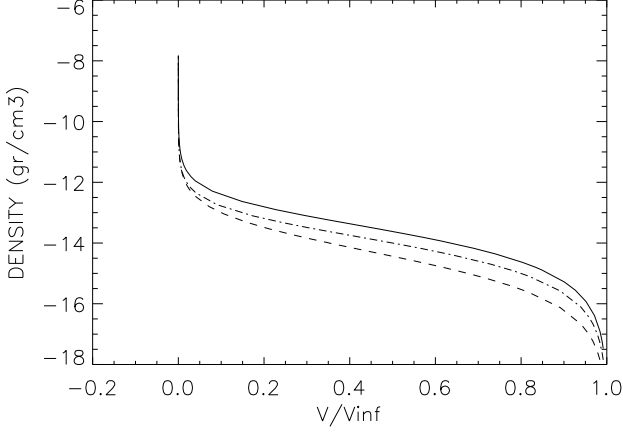


Fig. 4. Fit to the He II 4686 profile using different values for the terminal velocity and the β exponent of the wind velocity law. The mass-loss rate was also varied in each case to obtain a reasonable fit while keeping all other parameters fixed. Red, dotted line: $\beta = 2.0$, $V_\infty = 1800 \text{ km s}^{-1}$, $\dot{M} = 4.3 \times 10^{-7} M_\odot \text{ yr}^{-1}$ (adopted model); blue, dashed line: $\beta = 1.0$, $V_\infty = 1800 \text{ km s}^{-1}$, $\dot{M} = 1.1 \times 10^{-6} M_\odot \text{ yr}^{-1}$; orange, dot-dashed line: $\beta = 1.0$, $V_\infty = 900 \text{ km s}^{-1}$, $\dot{M} = 5.5 \times 10^{-7} M_\odot \text{ yr}^{-1}$. Black: Observed spectrum.

The value obtained ($\beta = 2.0$) is much higher than usual values for early and mid Of stars (0.7-1.0, see e.g. Repolust et al. 2004 for the MW or Mokiem et al. 2006 for the SMC). However, β and \dot{M} have different effects on the line profiles. A higher β means a slower acceleration in the inner wind (where the He II profile is formed) and therefore, because of the continuity equation ($\dot{M} = 4\pi r^2 \rho(r) v(r)$) for a fixed mass-loss rate and terminal wind velocity, at a given point r the velocity is lower and the density higher than for a smaller β (see Fig. 5). This implies more emission at low velocities, but also a lower velocity at a given density. The minimum density required to contribute to the emission profile is reached at lower velocities in the high β model. As a consequence, the line profile in this model has more emission at low velocities compared to the high β model. The result is a steeper red wing, fitted with a lower mass-loss rate in the high β model.

Table 2. Stellar parameters obtained here for GHV-62024 with FASTWIND. T_{eff} is given in K, g in cm s^{-2} , \dot{M} in solar masses per year and V_{∞} in km s^{-1}

Star ID	Sp Type	T_{eff}	$\log g$	R/R_{\odot}	$\log(L/L_{\odot})$	M_{sp}/M_{\odot}	\dot{M}	β	V_{∞}	Y_{He}	$\log(N/H) + 12$
62024	O6.5 IIIf	36500	3.60	11.1	5.29	18.3	4.30×10^{-7}	2.0	1800	0.18	8.30

**Fig. 5.** Run of the density versus wind velocity for three models. The first one (solid line) has $\dot{M} = 1.0 \times 10^{-6} M_{\odot} \text{yr}^{-1}$, $\beta = 1.0$ and the second one (dashed line) has $\dot{M} = 4.3 \times 10^{-7} M_{\odot} \text{yr}^{-1}$, $\beta = 2.0$. In both cases, the mass-loss rate is the one needed to fit the emission peak of He II. The third model (dot-dashed line) has $\dot{M} = 4.3 \times 10^{-7} M_{\odot} \text{yr}^{-1}$, $\beta = 1.0$. All other parameters are the same as in our final model.

In a preliminary analysis of GHV-62024, Herrero et al. (2011) assumed a standard value of $\beta = 0.9$ for its wind. As a result, the MWM rate of this star was much higher than predicted by the Vink et al. (2001) relationship for its luminosity and metallicity (note that if the star is contaminated by a companion, which is always a worry at far distances, the actual luminosity would be lower and thus the discrepancy with the theory larger). As indicated in the introduction, this result is similar to that found by Tramper et al. (2011) for stars in IC 1613, WLM and NGC 3109, who also adopted β (with values from 0.80 to 0.95, depending on luminosity class). As surprising as the β value obtained in this work may look, it solves the problem of the high MWM of GHV-62024, because it reduces the mass-loss rate while keeping all other parameters (nearly) fixed. Figure 6 illustrates this point. In the figure we plot the present $\log(D_{\text{mom}})$ value (28.22) and that obtained by Herrero et al. (2011) (28.69), together with the values given by Mokiem et al. (2007b) for Galactic and SMC stars. The stellar luminosity has also changed because we derive slightly different stellar parameters as a consequence of the new β . The new value for the MWM rate fits the canonical relationship defined by the SMC stars for unclumped winds. We note, however, that because the metallicity of IC 1613 is slightly lower than that of the SMC, we could expect the $\log(D_{\text{mom}})$ of GHV-62024 to be slightly lower as well.

The high β value obtained here is unusual, but there are precedents in the literature. The most similar case is probably that of AzV 83, analysed by Hillier et al. (2003). This is an O7 If star (as compared to O6.5 IIIf for GHV-62024) for which the authors also found a velocity law exponent $\beta = 2.0$. We do not include clumping in our analysis, but Hillier et al. did, therefore this is not the cause of the large β . This high value of β for such

an early star led the authors to suggest that AzV 83 is a fast rotator seen pole-on. The peculiar value for β would be the result of a mixture of photons from the asymmetric wind². Changing acceleration conditions along the wind may mimic a high β value (Owoc, private communication). Extending this argument, we may expect that the equatorial and polar winds are subject to different acceleration conditions, and accordingly that the global profile mimics a velocity law with a high β .

There are more similarities between the two stars. Hillier et al. (2003) obtained $Y_{\text{He}} = 0.20$ and $12 + \log(N/H) = 8.41$, compared to our $Y_{\text{He}} = 0.18$ and $12 + \log(N/H) = 8.30$. In both cases, the high abundances could be attributed to the proposed fast rotation. The main difference between both objects lies in the wind terminal velocity. AzV 83 has a modest $V_{\infty} = 940 \text{ km s}^{-1}$ as derived from its UV spectrum. This value is much lower than the 1800 km s^{-1} we estimated for GHV-62024 by scaling its escape velocity. The difference would have been smaller had we also used the escape velocity to determine the terminal wind velocity of AzV83. The escape velocities of the two stars are not so different: 436 km s^{-1} for AzV83 and 666 km s^{-1} for GHV-62024 (once more, it is true that the devil is in the details). Thus, when estimated from the respective escape velocities, the terminal velocities of both stars would differ by only $\sim 30\%$.

Could it be that the high β value derived here for GHV-62024 is the result of a too high estimation for V_{∞} ? As was shown by Puls et al. (1996), it is primarily the Q-parameter ($Q = \frac{\dot{M}}{(V_{\infty} R_{\star})^2}$) that determines the net wind emission. Therefore, for a given wind emission, if V_{∞} increases \dot{M} has to increase as well. And given the behaviour of density and velocity explained above and illustrated in Fig. 5, this could result in too much emission at high velocities. Could a lower V_{∞} result in a lower β value? The answer is no (see Fig. 4). Although a given stellar parameter has an effect on the whole profile of a given spectral line, its impact on different parts of the line is also different. Consequently β primarily determines the steepness of the red wing emission (because a higher β means a lower wind acceleration, and therefore the inner layers are denser and slower, resulting in more emission at a given wavelength but extending to lower velocities). On the other hand, the main impact of V_{∞} is in the extension of the lines profiles, primarily in the absorption part of the P-Cygni profile. In Fig. 4 we show an example with different values of V_{∞} and β . We see that we can reduce the red emission by decreasing the terminal velocity at $\beta = 1.0$, but then neither the slope of the red wing nor the extension of the blue absorption would fit. We conclude that our estimation of the wind terminal velocity cannot be wrong by more than 30%, which is consistent with the lower limit indicated in Sect. 3, and that this would not affect the derived β significantly.

² Hillier et al. (2003) adopted the equatorially compressed wind model from Bjorkman & Cassinelli (1993), i.e., the equatorial wind is denser than the polar one, in contrast to current views, which favour a faster and denser wind from the pole and a slower, thinner wind from regions closer to the equator. However, this difference is not important for the argument presented here.

AzV 83 is one of the stars considered by Mokiem et al. (2007b) (who adopted the values given by Hillier et al. 2003 after correcting for the clumping adopted by these authors) and appears in Fig. 6, in a position fully compatible with the rest of the stars. Again, a more standard β value would result in a MWM value higher than the average relation.

However, a too high wind terminal velocity would have a clear impact on $\log(D_{\text{mom}})$. As $\log(D_{\text{mom}}) = \log(\dot{M}V_{\infty}R^{0.5})$, using the above definition of Q we obtain $\log(D_{\text{mom}}) = 2.5 \log V_{\infty} + 2 \log R + \log Q$. For a fixed pair (Q, R) we obtain that a change of 30% in V_{∞} results in a change of 0.3 dex for $\log(D_{\text{mom}})$. This would bring the point of GHV-62024 in Fig. 6 close to the value expected for the IC 1613 relationship. This is what one would expect even for a fast rotator, because Maeder (2001) has shown that the WLR (and in turn the position of the star in Fig. 6) is not altered by fast rotation.

Therefore, we conclude that these two stars are consistent with a scenario in which they are fast rotators seen pole-on, their wind momentum is what is expected from their luminosities and metallicities. We suggest as a working hypothesis for future investigations that the fact that their β values are higher than usual may be a consequence of fast rotation.

Can we also observe higher β values for known fast rotators? Two fast rotators in the sample of Mokiem et al. (2007b) lie at positions close to GHV-62024 in the WLR diagram. NGC346-010 is an O7 III_{ne}(f) star analysed by Mokiem et al. (2006). NGC 346-010 is in fact quite similar to GHV-62024, except for the higher mass-loss rate and He abundance of the latter. Fig. 7 offers a comparison of the spectra of NGC346-010 and GHV-62024, with the last one convolved with an additional rotational profile with a velocity of 200 km s⁻¹ (to simulate the high rotational velocity of NGC346-010, but taking the much lower spectral resolution for GHV-62024 into account). In the figure we appreciate that both spectra are very similar, except for the emission at He II 4686 and N III 4630–40, a higher ratio He II 4542 / He I 4471 and the nebular emission in H_γ. The similarity in the rest of the spectral features indicates that the GHV-62024 emission in He II 4686 is a real feature. This figure can be compared with the spectrum of AzV 83, shown by Hillier et al. (2003) in their Fig. 2. There we see the strong emission in N 4634–40 and He 4686 (this time the line is fully in emission). Thus, the three stars look quite similar, except for these emission features (and the small differences in the other He profiles related to effective temperature, which could also be affected by differences in atmospheric extension). Because of the lack of UV data, the wind terminal velocity of NGC346-010 has been again estimated from the escape velocity (with $V_{\infty} = 1832$ km s⁻¹ in Mokiem et al. 2006, later revised to 1486 km s⁻¹ in Mokiem et al. 2007b). A standard value, $\beta = 0.8$, was adopted for the exponent of the velocity law. For this fast rotator, the authors suggest the presence of an equatorial disk (because of some difficulties in the line profile fitting, but without other evidence) although on the whole the fit to the H/He optical spectrum is satisfactory. The second fast rotator in the sample of Mokiem et al. (2007b) is AzV 296, a fast rotating O7.5 V((f)) analysed by Massey et al. (2004). In this case, the terminal velocity has been obtained from UV spectra ($V_{\infty} = 2000$ km s⁻¹). A $\beta = 0.8$ has also been adopted for the exponent of the velocity law. In this case the optical fit is poor and the authors suggest that the star might be a binary, which would also explain the relatively large radius needed to fit the observed magnitude and the high radial velocity observed.

In both cases, a $\log(D_{\text{mom}})$ slightly larger than the canonical relationship is derived. Unfortunately, all lines are strong in absorption including H_α, and therefore the β exponent could not

be accurately constrained. Consequently, the mass loss rate in this case may actually be an upper limit. A higher β value would therefore bring these two fast rotators into better agreement with the average WLR for their metallicity.

Therefore, we conclude that the analysis of these stars indicates that the velocity law may be altered in some stars. Anticipating data from the next section (where we discuss the N and He abundances), we suggest that fast rotation is a probable cause for this alteration. Note, however, that the evidence is indirect and based on very few objects. Moreover, of three stars with known “f” signature at low Z (AzV 83 and NGC 346-001 in the SMC and GHV-62024 in IC 1613), two have been suggested to be seen pole-on. This is quite improbable, unless the f-phenomenon at low Z is favoured by this pole-on view. If that were the case, the enhanced polar wind would be responsible for the presence of the f characteristics. This possibility agrees well with the findings by Rivero-González et al. (2011a), who found that the f-phenomenon is related to wind strength and might therefore decrease at low metallicity.

Mokiem et al. (2007b) have analysed other fast rotators in the SMC, but they give no further clues because the authors could only derive upper limits that are above the mean WLR for the SMC. However, it is interesting to note that we do not find this situation in the Milky Way: the values derived by Mokiem et al. (2007b) and Repolust et al. (2004), whether actual values or upper limits, are close to or below the average WLR for the Milky Way and we are not aware of significantly high β values needed to fit the spectra of Galactic O stars, although they are found for the slower winds of B-supergiants (Kudritzki et al. 1999; Crowther et al. 2006; Markova & Puls 2008). We note that determining β from optical spectra of rapidly rotating stars seen edge-on is in most cases very difficult because of the lack of spectral features sensitive to this parameter.

We have not taken clumping into account in our analysis and therefore the derived mass-loss rate has to be considered as an upper limit (Puls, Vink & Najarro 2008). However, as far as we compared ours with other results and theoretical predictions that also use unclumped winds, our conclusions remain valid. Of course, there might be relative differences in the clumping among stars with different properties, in particular the winds of fast rotators might be clumped in a different way than those of slowly rotating stars. However, studies about the clumping distribution are at their very beginning, and therefore we have no data yet to discuss this point.

6. The evolutionary status of GHV-62024

In Fig. 8 we place GHV-62024 in the Hertzsprung-Russell diagram and compare it with evolutionary tracks from Brott et al. (2011) for the SMC. We chose the tracks with an initial rotational velocity $V_{\text{rot}} = 410$ km s⁻¹ because they have the highest rotational velocity consistent with the star locus. For higher rotational velocities, the corresponding tracks evolve bluewards already from the ZAMS.

We see that according to these tracks the star was born with an initial mass of $\sim 30 M_{\odot}$. This result does not change significantly when using other tracks, even those without initial rotational velocity. The present evolutionary mass derived from the adopted tracks (at $V_{\text{rot,ini}} = 410$ km s⁻¹) is $28.6 M_{\odot}$. A comparison with the spectroscopic mass reveals then the well-known mass-discrepancy.

The most interesting parameter for the evolutionary status of GHV-62024 is its nitrogen abundance. We obtained a value of $\log(N/H) + 12 = 8.30 \pm 0.30$ dex from our analysis in Sect. 4.

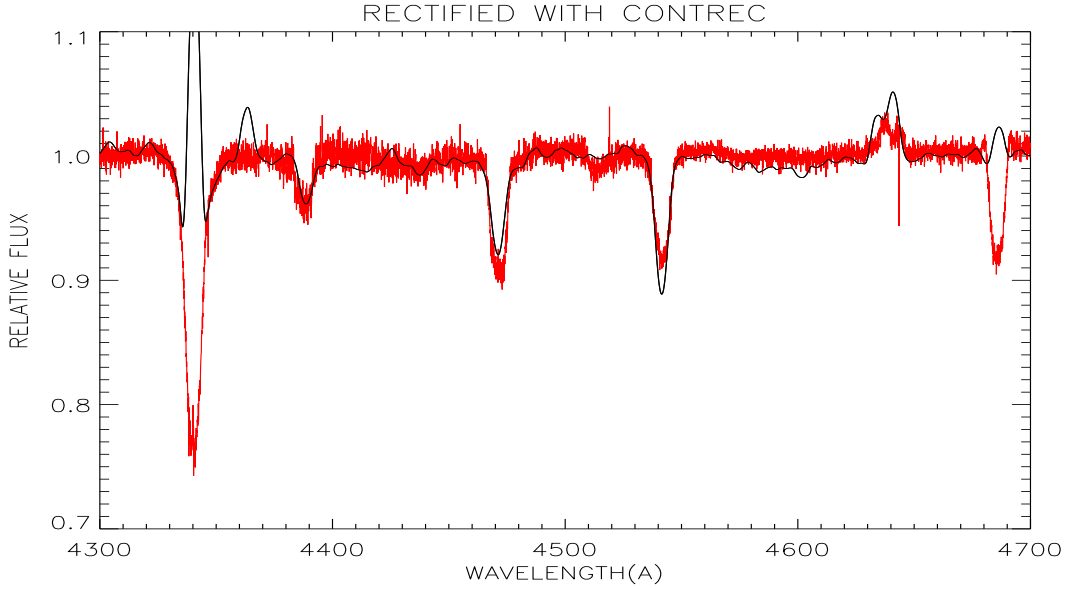


Fig. 7. Comparison of the spectra of GHV-62024 (black) and NGC346-010 (red).

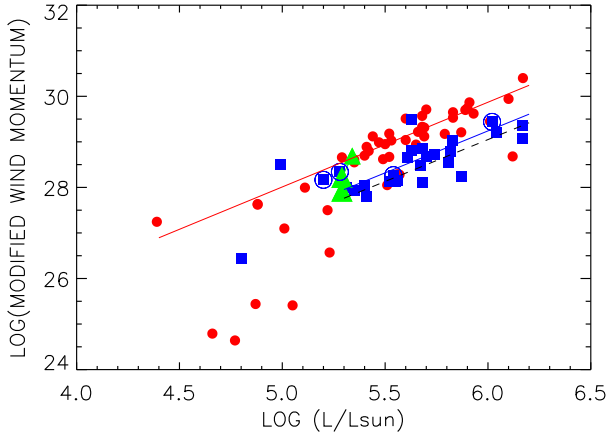


Fig. 6. Wind-momentum luminosity relationship for GHV-62024 (upper green triangle: value from Herrero et al. (2011); middle green triangle: this work; lower green triangle: this work, considering a wind terminal velocity 30% lower than quoted in Tab. 2) compared to data from Mokiem et al. (2007b) for the MW (red circles) and SMC (blue squares). The SMC stars cited in text were marked with an open circle. From left to right they are NGC346-010, AzV296, AzV83 and NGC346-001. The solid lines represent the WLR fitted by Mokiem et al. (2007b) without considering clumping for MW and SMC objects, and the black dashed line corresponds to the SMC relation shifted to $Z = 0.13Z_{\odot}$.

This strong nitrogen enhancement (and the similarly high He abundance, $Y_{\text{He}} = 0.18 \pm 0.04$) cannot be explained with non-rotating models, which only predict N and He enhancement much later, in the red supergiant phase. Therefore, we looked at the set of rotating models calculated by Brott et al. (2011) for the SMC metallicity. Fig. 9 presents the predictions of models with $V_{\text{rot}} = 410 \text{ km s}^{-1}$ compared with the value derived from the observations. The models predict a slightly lower maximum N abundance (except for homogeneous evolution) and never

reach the derived value of 8.3, although they are consistent with the lower half of the error bar. We have to take into account that the highest possible N enrichment is limited by the initial abundances of C and O (which decrease with the metallicity), and that the C and O abundances could scale differently from solar; alternatively, there might still be some problems in our N III line-synthesis (Rivero-González et al. 2011a), which might add to the quoted error budget. The He abundance is also higher than predicted for a model with an initial mass $M_0 = 30 M_{\odot}$ at the chosen initial rotational velocity, but would be consistent with that of a model of $40 M_{\odot}$, for which $Y_{\text{He}} = 0.20$ is predicted at the temperature obtained for GHV-62024, or with that of a slightly faster rotating model.

Our results qualitatively agree with those from Rivero-González et al. (2011b) for the LMC, who also obtained very high N and He abundances for a large fraction of their LMC sample. GHV-62024 would very nicely fit in Rivero-González et al. Fig. 8 (lower panel), where they show a clear correlation between the abundances of these two elements. This indicates that in both cases the N and He enrichment has the same origin, presumably the exposition at the surface of the products of the CNO bi-cycle taking place in the stellar interior.

Using the same tracks, we can estimate the stellar age to be less than 4.5 Myr, in qualitative agreement with our expectations for a young region. In Fig. 10 we plot the behaviour of the rotational velocity versus the effective temperature in the models. Interestingly, we see that the rotational velocity is predicted to remain at high values until much later in the evolution of the star (the effective temperature acts here as a temporal variable). This is fully consistent with our expectation that the star is a fast rotator. Therefore, we think that it is possible to conclude that GHV-62024 shows strong N and He abundances that are compatible with a high initial rotational velocity still present in the star, although other possibilities (such as an interactive binary history) cannot be discarded with the present data.

However, if the star still retains its high rotational velocity (note in Fig. 10 that for the model with $30 M_{\odot}$ the rotational velocity increases during the early evolution), this could mean that 2D effects may become important. In that case, our 1D models would be too simplified and orientation effects could play a role.

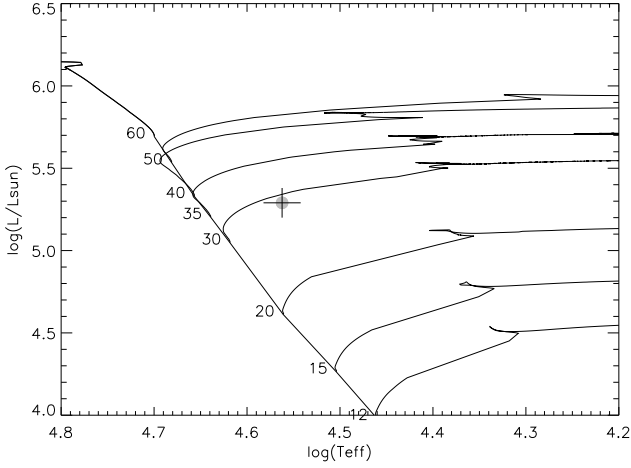


Fig. 8. Position of GHV-62024 in the HR Diagram. Evolutionary tracks are taken from Brott et al. (2011), with an initial rotational velocity of 410 km s^{-1} .

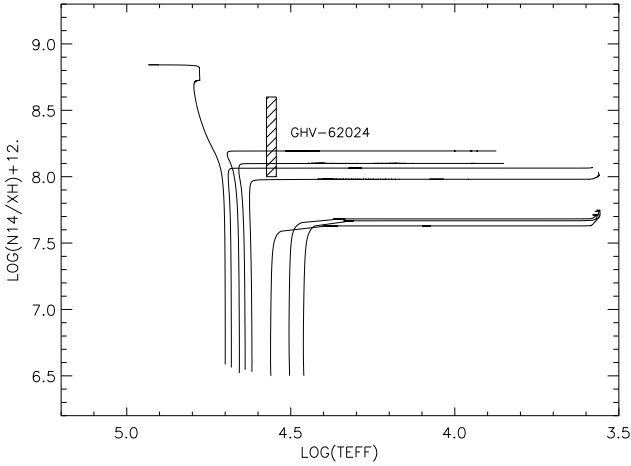


Fig. 9. Variation of the abundance of ^{14}N along the stellar lifetime (as given by the evolution in temperature). Represented in the ordinate axis is $\log(^{14}\text{N}/\text{H}) + 12$. Models with a zero initial rotational velocity only change their nitrogen abundance in the red supergiant phase, whereas models with an initial rotational velocity of 410 km s^{-1} suffer an early nitrogen enhancement. Evolutionary models have masses of 60, 50, 40, 35, 30, 20, 15 and $12 M_{\odot}$, from left to right. The position of GHV-62024 is represented by the dashed box.

A proper comparison would have been made between 2D evolutionary and atmospheric models, which is a work for the future. However, we are confident that even if the high rotational velocity is confirmed, the analysis and comparisons described here are a good representation of the average properties of the star.

7. Conclusions

We have analysed GHV-62024, an O6.5 IIIf star in the low-metallicity galaxy IC 1613.

Our preliminary analysis of this star (Herrero et al. 2011) resulted in a too high MWM rate for the star, typical of Milky Way metallicity stars. This would have posed a problem for the

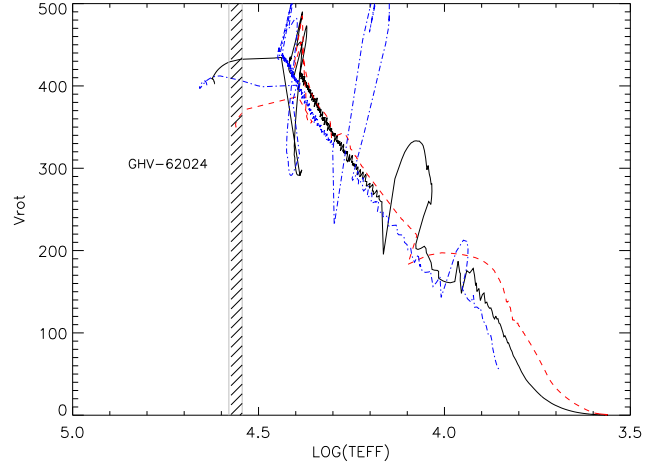


Fig. 10. Evolution of the rotational velocity for evolutionary models with initial masses of 30 (black), 35 (blue, dot-dashed line) and 20 (red, dashed line) solar masses, according to Brott et al. (2011). The temperature range derived for GHV-62024 is indicated by the hatched region.

theory of radiatively driven winds. However, in our new analysis presented in this work we conclude that the MWM is consistent with a lower-than-SMC metallicity, but we derive a β exponent of 2.0 for the wind velocity law. This is higher than the usual value for stars of this spectral type. Following a suggestion by Hillier et al. (2003), we hypothesize that the star is a fast rotator seen from low inclination angle and that the high β would then be an artefact caused by the superposition of light emitted from regions with different wind properties.

When comparing stellar evolutionary models with rotation from Brott et al. (2011) we see that the star shows the well-known mass-discrepancy. We derive high N ($\log(\text{N}/\text{H}) + 12 = 8.30 \pm 0.30$) and He ($Y = 0.18 \pm 0.04$) abundances, which can qualitatively be explained by a strong mixing with processed CNO products. The N abundance is higher than the maximum value predicted by the tracks from Brott et al. (2011) for non-homogeneous evolution, although it is consistent within the errors, and the He abundance is consistent with the model prediction either for a higher mass than the initial stellar mass derived from the position of the star in the HRD or for a faster initial rotational velocity. However, with an initial rotation of 410 km s^{-1} the models predict that the star can reach high abundances while still keeping the high rotational velocity. This is therefore consistent with the above hypothesis that GHV-62024 is a fast rotator seen from a low inclination angle, whose velocity field properties would have been modified towards high β values without modifying the average WLR (in agreement with Maeder 2001 and Petrenz & Puls 2000).

However, a high rotational velocity could give rise to 2D effects, not included in the atmosphere or evolutionary models. If the high rotational velocity is confirmed, the analysis and comparisons carried out in this work should be considered as a first step towards a more complete, multidimensional analysis.

Acknowledgements. We want to thank an anonymous referee for very useful comments that helped to improve our paper. This work was supported by the Spanish Ministerio de Ciencia e Innovación (grants AYA2008-06166-C03-01 and -02, AYA2010-21697-C05-01 and -04) and the Consolider-Ingenio 2010 Program (CSD2006-00070) and the Gobierno de Canarias (grant PID2010119). J.G.R.G. gratefully acknowledges financial support from the German DFG, un-

der grant 418 SPA 112/1/08 (agreement between the DFG and the Instituto de Astrofísica de Canarias).

References

- Asplund, M., Grevesse, N., Sauval, Scott, P., 2009, *ARA&A*, 47, 481
- Bjorkman, J.E., & Cassinelli, J.E., *ApJ*, 409, 429
- Bouret, J.C., Lanz, T., Hillier, D.J. et al., 2003, *ApJ*, 595, 1182
- Bresolin, F., Kudritzki, R.P., Mendez, R.H. & Przybilla, N., 2001, *ApJ*, 548, 159
- Bresolin, F., Urbaneja, M.A., Gieren, W., Pietrzyński, G. & Kudritzki, R.P., 2007, *ApJ*, 671, 2028
- Brott, I., de Mink, S., Cantiello, M. et al., 2011, *A&A*, 530, A115
- Castor, J.I., Abbott, D.C. & Klein, R.I., 1975, *ApJ*, 170, 325
- Castro, N., 2010, PhD, Univ. La Laguna
- Castro, N., Urbaneja, M.A., Herrero, A., et al., 2012, *A&A*, in press
- Conti, P.S. & Alschuler, W.R., 1971, *ApJ*, 170, 325
- Crowther, P.A., Lennon, D.J. & Walborn, N.R., 2006, *A&A*, 446, 279
- Dolphin, A.E., Saha, A., Skillman, E.D. et al., 2001, *ApJ*, 550, 554
- Dufton, P.L., Ryans, R.S.I., Trundle, C. et al., 2005, *A&A*, 434, 1125
- Evans, C.J., Smartt, S.J., Lee, J.-K. et al., 2005, *A&A*, 437, 467
- Evans, C.J., Bresolin, F., Urbaneja, M.A. et al., 2007, *ApJ*, 659, 1198
- Evans, C.J., Taylor, W.D., Hénault-Brunet, V. et al., 2011, *A&A*, 530, A108
- Friend, D.B. & Abbott, D.C., 1986, *ApJ*, 311, 701
- García, M., Herrero, A., Vicente, B. et al., 2009a, *A&A*, 502, 1015
- García, M., Herrero, A., Castro, N., Corral, L., Rosenberg, A., 2010, *A&A*, 523, A23
- Heap, S.R., Lanz, T. & Hubeny, I., 2006, *ApJ*, 638, 409
- Herrero, A., Kudritzki, R.P., Vilchez, J.M. et al., 1992, *A&A* 261, 209
- Herrero, A., García, M., Uytterhoeven, K. et al., 2010, *A&A*, 513, A70
- Herrero, A., García, M., Uytterhoeven, K. et al., 2011, *IAUS* 272, 292
- Hillier, D.J., Lanz, T., Heap, S.R. et al., 2003, *ApJ*, 588, 1039
- Hunter, I., Brott, I., Lennon, D.J. et al., 2008, *ApJ*, 676, L29
- Kudritzki, R.P., Puls, J., Lennon, D.J. et al., 1999, *A&A*, 350, 970
- Kudritzki, R.P. & Puls, J., 2000, *ARA&A*, 38, 613
- Lee, H., Grebel, E.K. & Hodge, P.W., 2003, *A&A*, 401, 141
- Leitherer, C., Robert, C. & Drissen, L., 1992, *ApJ* 401, 596
- Lu, N.Y., Hoffman, G.L., Groff, T., Roos, T. & Lamphier, C., 1993, *ApJS*, 88, 383
- Maeder, A., 2001, *A&A*, 373, 122
- Markova, N. & Puls, J., 2008, *A&A*, 478, 823
- Markova, N., Puls, J., Scuderi, S., Simón-Díaz, S., Herrero, A., 2011, *A&A*, 530, A11
- Massey, P., Bresolin, F., Kudritzki, R.P., Puls, J., Pauldrach, A.W.A., 2004, *ApJ*, 608, 1001
- Massey, P., Puls, J., Pauldrach, A.W.A. et al., 2005, *ApJ*, 627, 477
- Mokiem, M.R., de Koter, A., Evans, C.J. et al., 2006, *A&A*, 456, 1131
- Mokiem, M.R., de Koter, A., Evans, C.J. et al., 2007, *A&A*, 465, 1003
- Mokiem, M.R., de Koter, A., Vink, J.S. et al., 2007, *A&A*, 473, 603
- Pauldrach, A., Puls, J. & Kudritzki, R.P., 1986, 164, 86
- Petrenz, P. & Puls, J., 2000, *A&A*, 358, 956
- Puls, J., Kudritzki, R.P., Herrero, A. et al., 1996, *A&A*, 305, 171
- Puls, J., Vink, J.S., & Najarro, F., 2008, *A&ARv*, 16, 209
- Repolust, T., Puls, J. & Herrero, A., 2004, *A&A*, 415, 349
- Rivero-González, J., Puls, J. & Najarro, F., 2011a, *A&A*, 536, 58
- Rivero-González, J., Puls, J. & Najarro, F., & Brott, I., 2011b, *A&A*, in press, arXiv:1110.5148
- Tramper, F., Saha, H., de Koter, A. & Kaper, L., 2011, *ApJ*, 741, 8
- Urbaneja, M.A., Herrero, A., Bresolin, F. et al., 2005, *ApJ*, 622, 862
- Urbaneja, M.A., Kudritzki, R.P., Bresolin, F. et al., 2008, *ApJ*, 684, 118
- Vink, J.S., de Koter, A. & Lamers, H.J.G.L.M., 2001, *A&A*, 369, 574